Cosmological implications of photon-flux upper limits at ultra-high energies in scenarios of Planckian-interacting massive particles for DM

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Study motivation

- Various null results for WIMP searches: originally expected masses pushed towards larger values and couplings towards weaker ones
- Post-LHC era: quartic coupling of the Higgs never too negative up to the Planck mass to induce instability $\Longrightarrow$ SM may be extrapolated up to Planck mass without encountering any inconsistency
- Inflationary cosmologies: SHDM production during reheating possible through minimal coupling (gravitation)
- SHDM decay possible in minimal-coupling scenarios through non-perturbative effects
- By-product decays detectable in UHECR data

Naturalness, WIMPs and Dark Matter

☛ Particle physics

- Problem of the Higgs mass: as a scalar field, can be destabilized by one-loop radiative corrections through its coupling to the top quark (quadratic divergences)

$$\delta m_h^2 = \frac{3\Lambda^2}{8\pi^2 v^2}\left[(4m_t^2 - 2M_W^2 - M_Z^2 - m_h^2) + \log\left(\frac{\Lambda}{\mu}\right)\right]$$

- Naturalness: stability of observables should prevail under small variations of the fundamental (bare) parameters $\implies \delta m_h^2 < m_h^2 \implies \Lambda < 1$ TeV – scale of new physics
- Supersymmetry or extra dimensions: add through various mechanisms to the spectrum of elementary particles other ones, one of which would be stable with a mass around 100 GeV and weak couplings

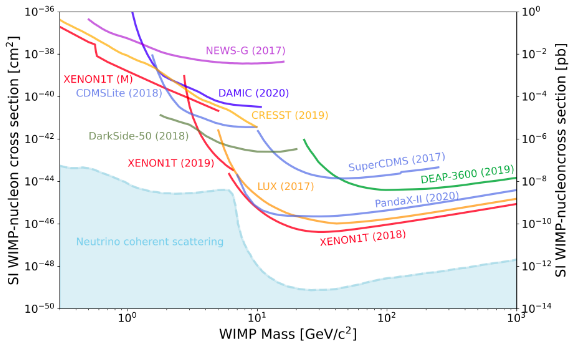
☛ Cosmology

- Freezing time estimated by equating the annihilation rate with the Hubble parameter $\implies \Omega_{\mathrm{WIMP}} \sim \frac{10^{-25}\ \mathrm{cm^3 s^{-1}}}{\langle \sigma v\rangle}$
- Of the order of unity by taking, as (should be) expected for WIMPs, $\langle \sigma v\rangle \sim G_F^2 M_X^2 \to$ the WIMP "miracle"

WIMPs?

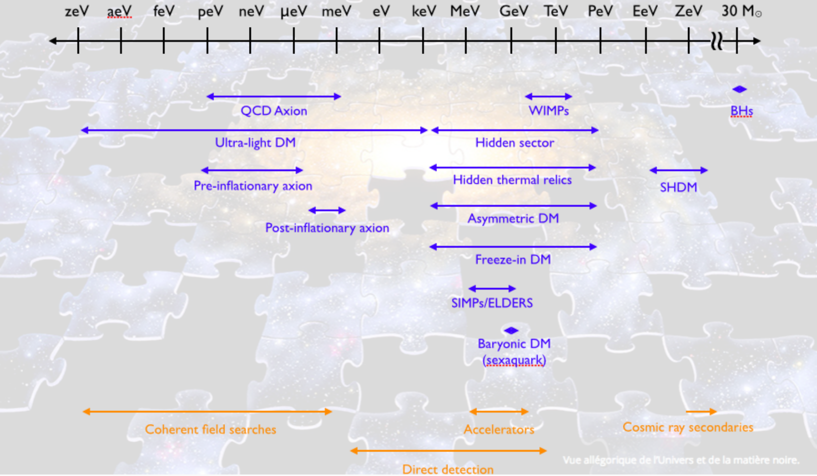
- Direct-detection searches: measurement of nuclear recoil

PDG, Prog. Theor. Exp. Phys. 2020 (2020) 083C01



- Accelerator-based and indirect-detection searches also unsuccessful

Dark Matter?



inspired from arXiv:1707.04591

SHDM motivations? SM vacuum (in)stability

☛ Alternative to naturalness to probe the energy scale Λ: SM vacuum (in)stability $\rightarrow$ very simplified calculation below, just a trend showing the necessity of new physics at scale Λ to avoid instability and the leading role of m_h and m_t

- To lowest order in the Higgs self-coupling λ, $\lambda(\mu)$ evolution dominated by the term from the top coupling (one-loop radiative correction):

$$\frac{\mu \mathrm{d}\lambda}{\mathrm{d}\mu} = -\frac{3\lambda_t^4}{8\pi^2} + \dots$$

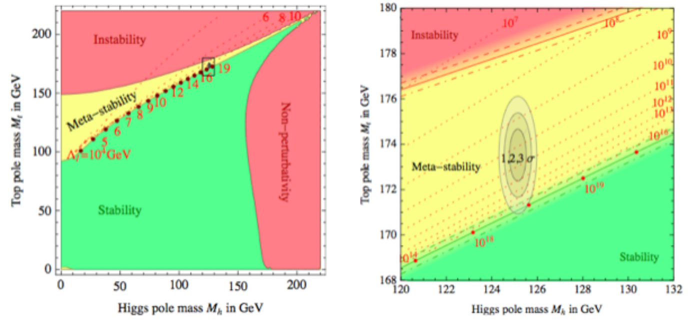
- As soon as $\lambda(\mu)$ turns negative, the Higgs potential becomes unbounded from below and the vacuum can suffer from instability
- Neglecting gauge interactions, the solution of the RGE at the instability scale $\lambda(\Lambda) = 0$ relates the Higgs mass with the top Yukawa coupling:

$$m_h^2 > \frac{3m_t^4}{\pi^2 v^2} \log \frac{\Lambda}{v}$$

LHC SM phase diagrams

- Extrapolation of the SM parameters up to large energies with full 3-loop NNLO precision

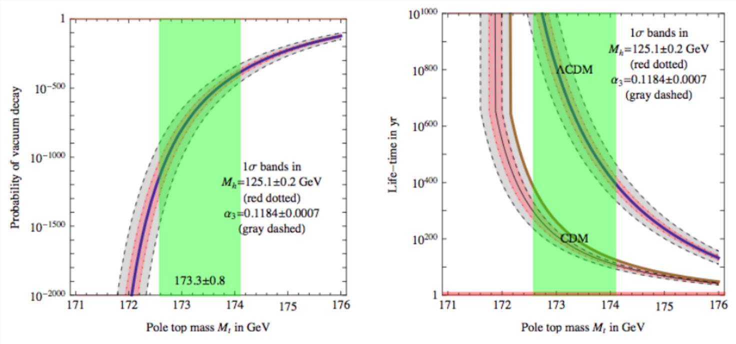
D. Buttazzo et al., JHEP vol 2013, 89 (2013)



- Precise values of Higgs boson mass + top Yukawa coupling $\Longrightarrow$ SM vacuum *meta-stable* up to high Λ

LHC SM phase diagrams

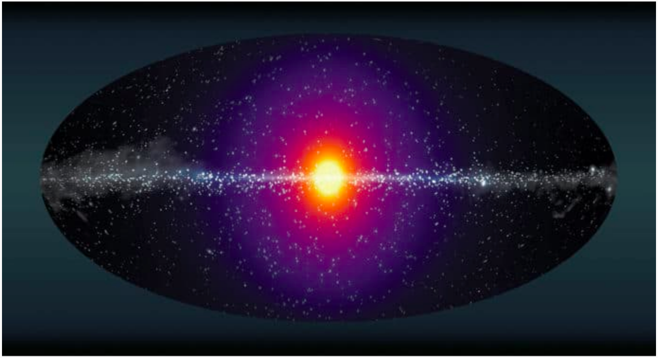
D. Buttazzo et al., JHEP vol 2013, 89 (2013)



- No inconsistency that would make the SM vacuum unstable by extrapolating the SM all the way from the mass of the top to the Planck mass
- Dark sector of super-heavy particles?

Dark-Matter Galaxy

- Rotation curves
- Large-scale structure simulations from primordial fluctuations
- Energy density $\rho_\odot = 0.3$ GeV cm^{-3}



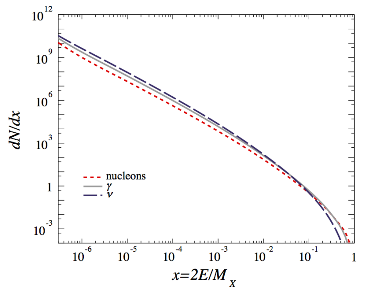
UHE particles as secondaries of SHDM decays

- For n pairs of $q\bar{q}$,

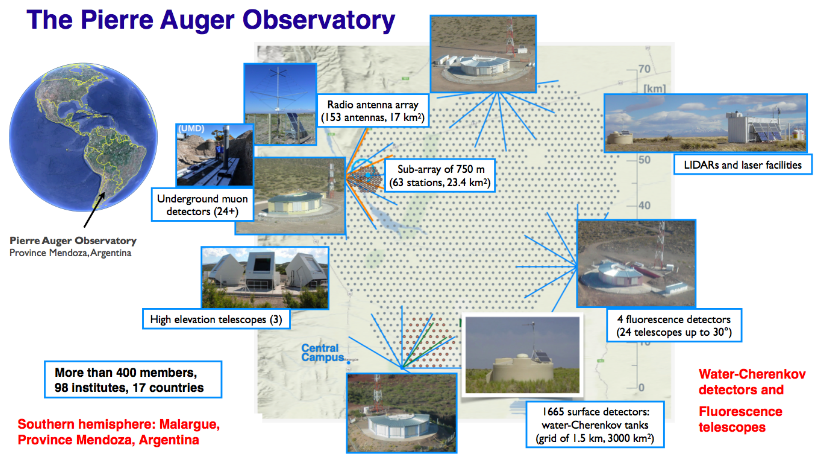
$$\frac{\mathrm{d}N_\gamma(x)}{\mathrm{d}x} = \frac{n(n-1)(n-2)\epsilon_\pi}{3} \int_x^1 \frac{\mathrm{d}z}{z} \frac{x}{z} \left(1 - \frac{x}{z}\right)^{n-3} \frac{D_h(z)}{z},$$

- ϵ_π: “efficiency” of the hadronization process into pions

- $D_h(z)$: fragmentation function of a parton into a hadron obtained from the fragmentation functions of partons evolved starting from measurements at the electroweak scale up to the energy scale fixed by M_X using the DGLAP equation

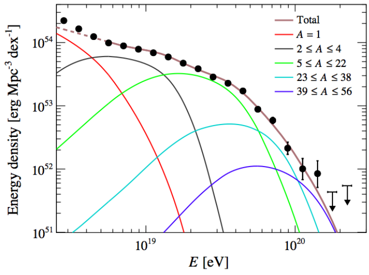


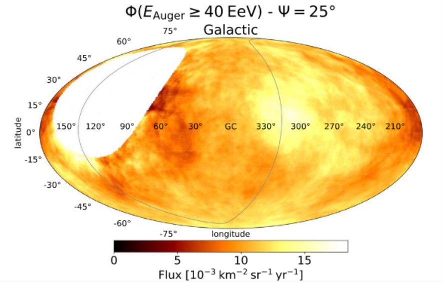
The Pierre Auger Observatory



UHE Cosmic Rays

- Nuclei from proton to CNO and Fe
- Accelerated in extragalactic sources



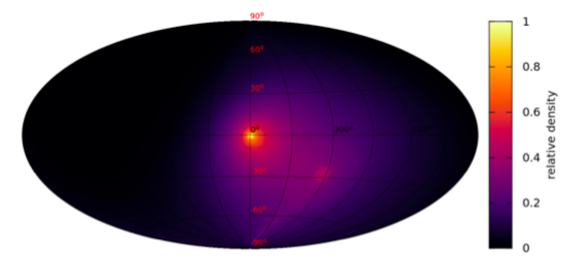


Anisotropy signatures of SHDM by-product decays

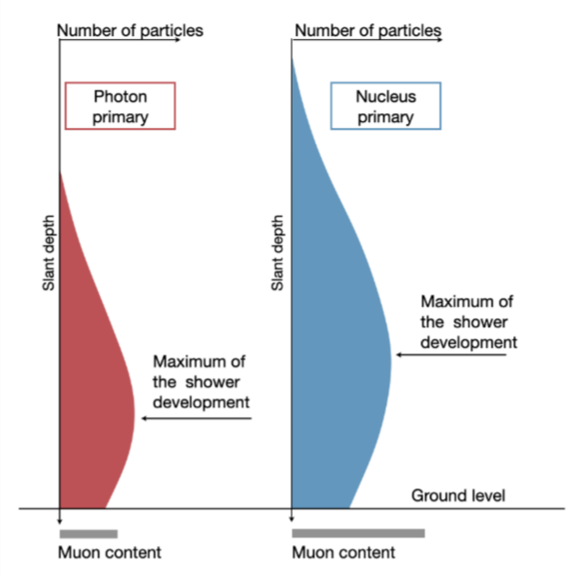
- Flux of secondaries from SHDM decay ($i = \gamma, \nu, \overline{\nu}, N, \overline{N}$):

$$J_i^{\mathrm{gal}}(E) = \frac{1}{4\pi M_X \tau_X} \frac{\mathrm{d}N_i}{\mathrm{d}E} \int_0^\infty \mathrm{d}s\; \rho_{\mathrm{DM}}(\mathbf{x}_\odot + \mathbf{x}_i(s;\mathbf{n})).$$

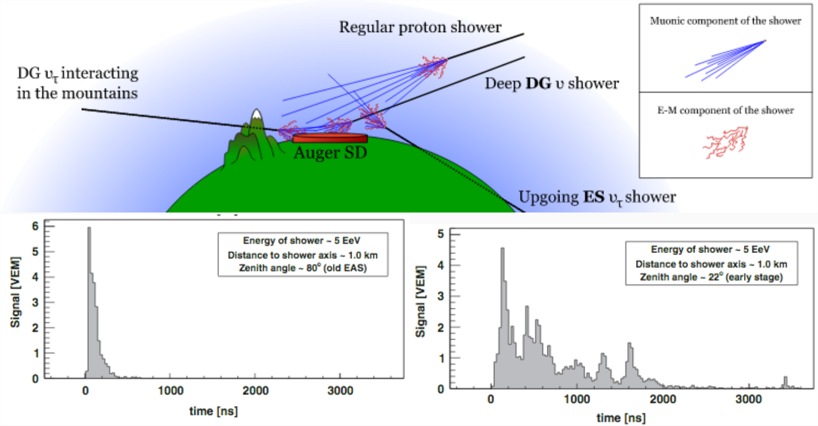
- ρ_{DM}: DM profile
- $\frac{\mathrm{d}N_i}{\mathrm{d}E}$: energy spectra of $i = \gamma, \nu, \overline{\nu}, N, \overline{N}$ from hadronization processes, evolving the fragmentation functions from EW scale up to M_X using DGLAP [Aloisio et al., Phys. Rev. D 69 094023 (2004)]
- Free parameters: M_X, τ_X



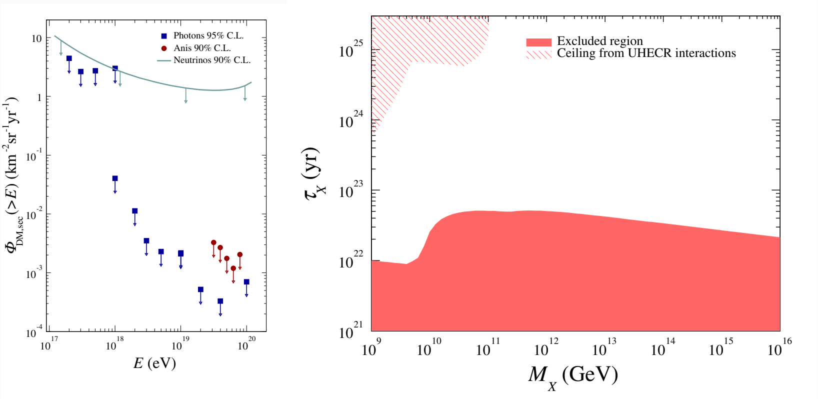
Searches for UHE photons



Searches for UHE neutrinos



Upper limits

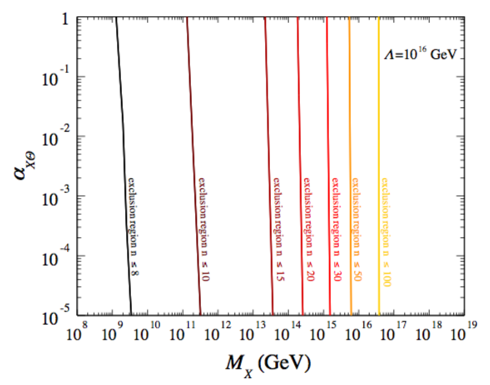


Constraints on perturbative decay

- Decay rate for an effective interaction term containing a monomial of dimension *n* in mass unit:

$$\Gamma_X \propto \alpha_{X\Theta} M_X \left(\frac{M_X}{\Lambda} \right)^{2n-8} .$$

- Fine tuning between $\alpha_{X\Theta}$ and *n*



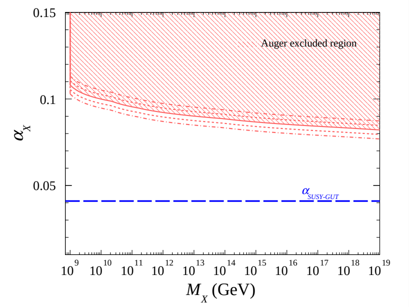
Constraints on non-perturbative decay

[Kuzmin & Rubakov, Phys. Atom. Nucl.979 61, 1028 (1998)]

- SHDM particles protected from standard decay by perturbative effects through a new quantum number
- Still, non-perturbative effects can lead to decays through “instantons” in non-commutative gauge theories
- For *B*, *L* and *X* currents not associated to gauge interactions, possibility to exchange quantum numbers through an anomaly

- Lifetime of metastable *X* particles: $\tau_X \simeq M_X^{-1} \exp(4\pi/\alpha_X)$

[t'Hooft, PRL 37 (1976) 8]



Non-thermal SHDM production during reheating

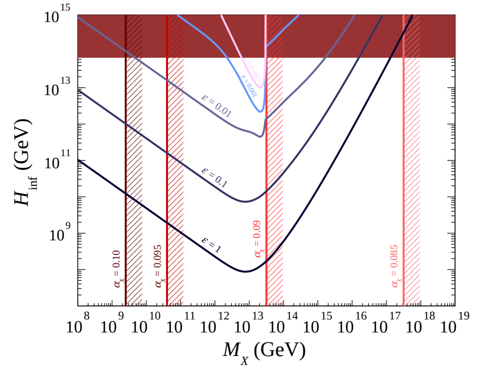
- No coupling between SM and DM sectors except gravitational
- DM production by "freeze-in" mechanism through s-channel SM+SM$\rightarrow$DM+DM [Garny et al. PRL 116 (2016) 101302] or $\phi + \phi \rightarrow$DM+DM [Mambrini & Olive Phys. Rev. D 103 (2021) 11, 115009] while inflaton decays into SM particles and reheats the universe after inflation:

$$\frac{dn_X(t)}{dt} + 3H(t)n_X(t) \simeq \sum_i \overline{n}_i^2 \Gamma_i$$

- Reheating dynamics between $t = H_{\rm inf}^{-1}$ and $t = \Gamma_\phi^{-1}$ at $T_{\rm rh}$ [Chung et al. Phys. Rev.929 D 60, 063504 (1999), Giudice et al., Phys. Rev. D 64, 023508 (2001)]:
    - $T(a) \simeq 0.2(\epsilon M_{\rm Pl} H_{\rm inf})^{1/2} \left(a^{-3/2} - a^{-4}\right)^{1/4}$
    - $H(a) = H_{\rm inf}(a/a_{\rm inf})^{-3/2}, a \leq a_{\rm rh}$
    - $H(a) = H_{\rm inf}\epsilon^2(a/a_{\rm rh})^{-2}, a > a_{\rm rh}$
- Reheating efficiency $\epsilon \simeq 4T_{\rm rh}(M_{\rm Pl}H_{\rm inf})^{-1/2}$ defined between 0 and 1, characterizing the duration of the reheating period ($\epsilon \simeq 1 \implies$ instantaneous reheating)

Viable regions

- Delineating viable regions in the $(H_{\rm inf}, M_X)$ plane for various ϵ values to match the DM relic density



- GUT mass scale viable for $\epsilon \to 1$ ($T_{\rm rh}$ relatively high) $\implies$ tensor/scalar ratio r of the primordial modes possibly detectable in the CMB
- For $\epsilon \leq 0.01$, 10^{13} GeV mass scale viable, testable for $\alpha_X \lesssim 0.09$

Summary

- Assuming no new physics up to high energy scales, several constraints on the properties of a dark sector of SH particles brought by the absence of UHE photons
- X particles with masses as large as the GUT energy scale could be sufficiently abundant to match the DM relic density, provided that the inflationary energy scale is high ($H_{\rm inf} \simeq 10^{13}$ GeV) and $T_{\rm rh}$ is high (so that reheating is quasi-instantaneous)
- UHECR/CMB complementarity